

# Dynamics of Polar Jets from the Chromosphere to the Corona: Mass, Momentum and Energy Transfer

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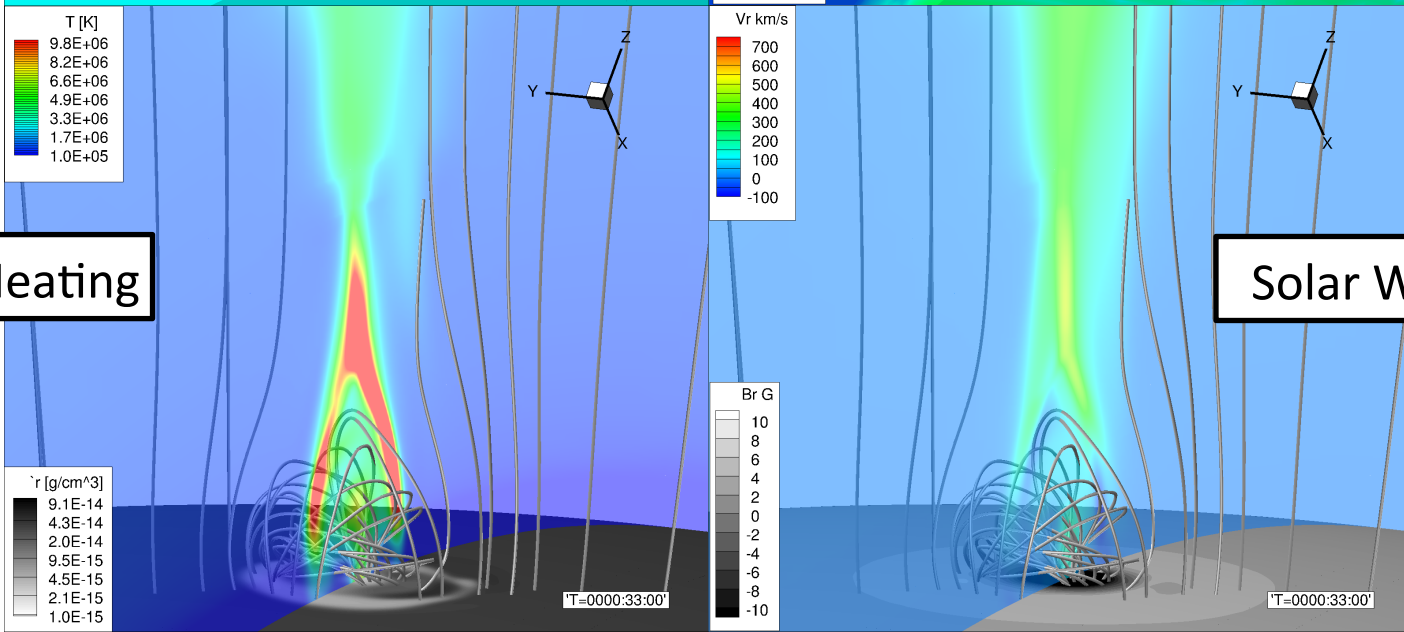
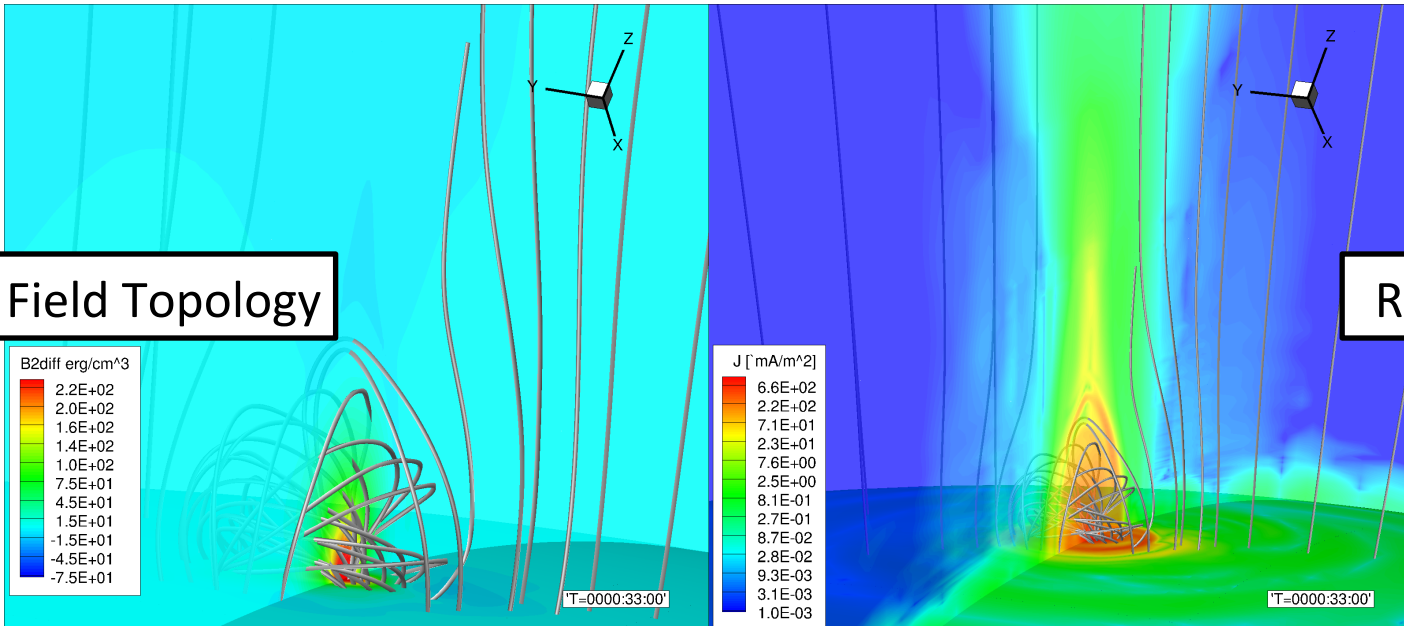
SPA-Solar and Heliospheric Physics

Tuesday, December 15, 2015

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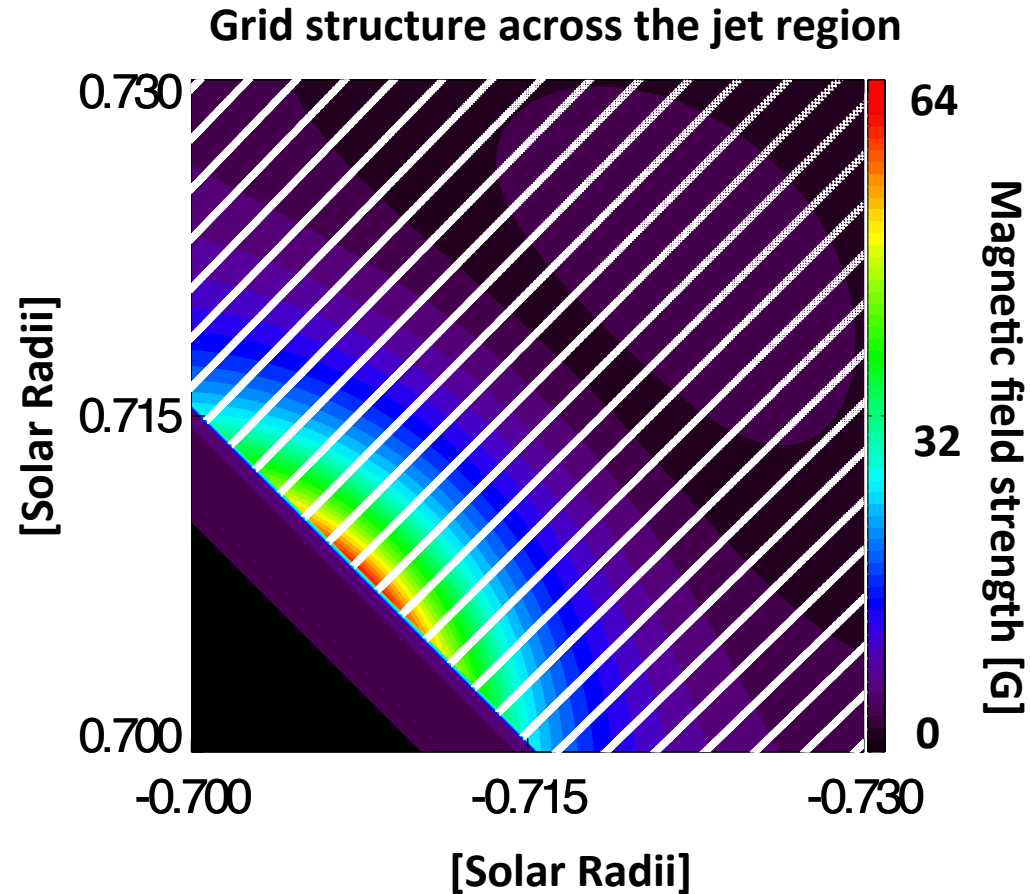


# Why Coronal Jets are Interesting?



# Goal of this study

- Study transport due to jet
  - energy, momentum, mass
  - interaction with local plasma
- State of the art model
  - physically well-established two temperature coronal model: AWM
  - full 3D MHD corona ( $24 R_{\odot}$ )
  - fully resolved small-scale jet structure ( $0.013 R_{\odot}$ )

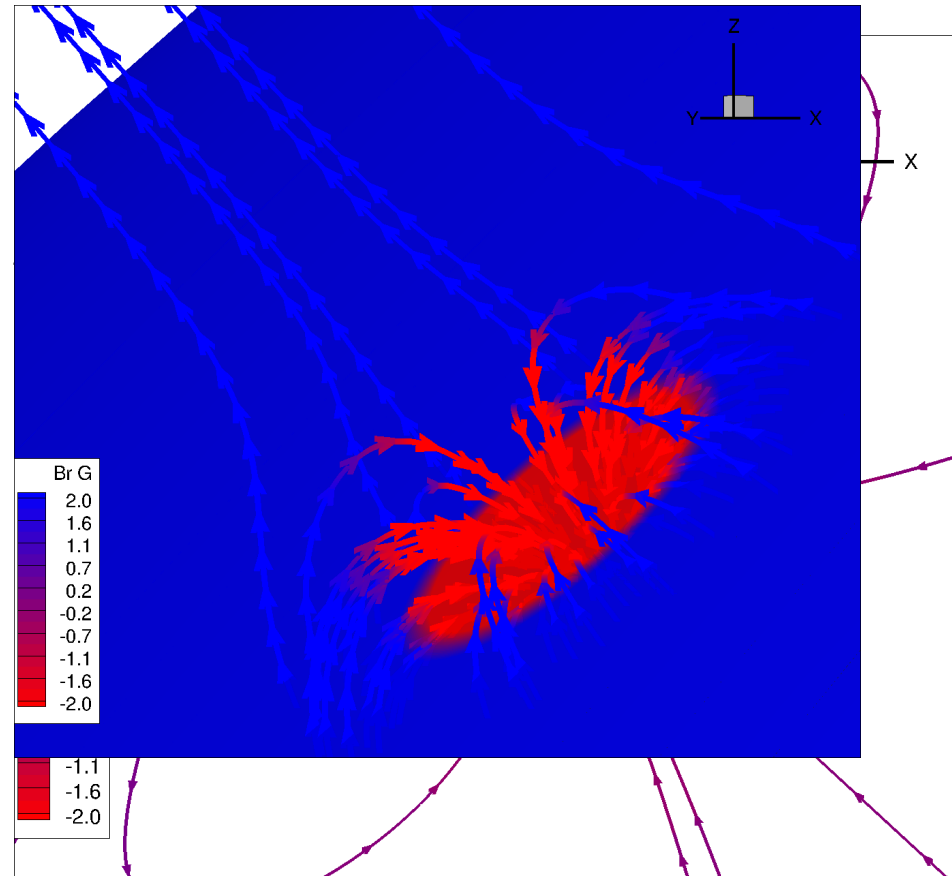


**Cell size: 1000 km in azimuthal,  
20 km in radial direction**

# Model Setup

- GLOBAL: Solar wind model
  - Chromospheric inner-boundary
  - Ambient dipole: -5.6 G at pole
  - Coronal heating and solar wind acceleration via Alfvén waves
- LOCAL: Jet model
  - Radial bipole under the surface
  - Rotating boundary condition around bipole axis
  - Magnetic latitude determines background field: open vs. closed

This talk focuses on polar jets.



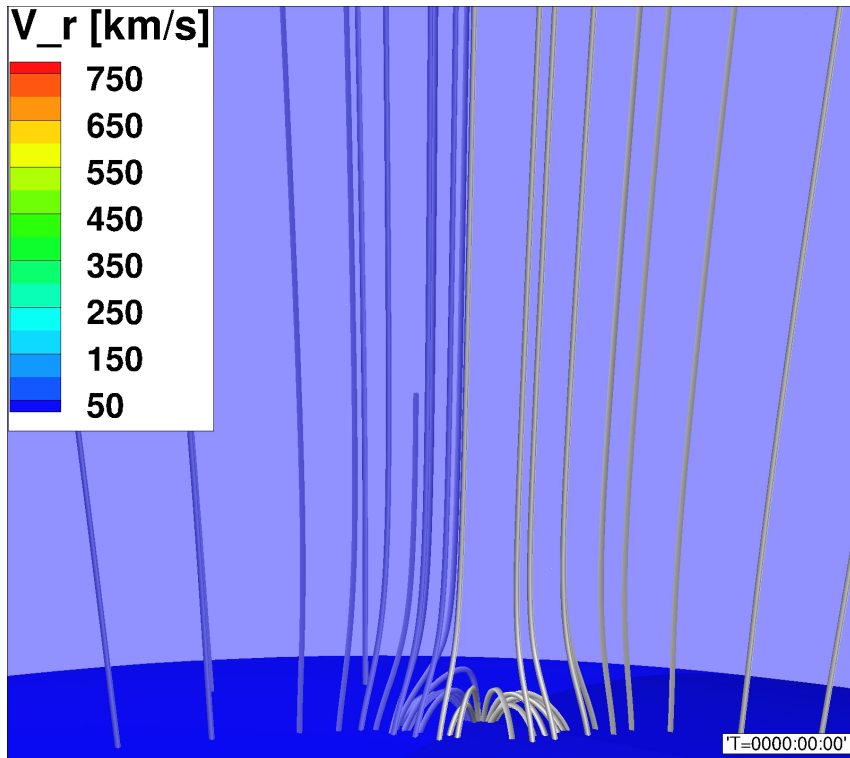
## Jet Model Parameters

Strength at 1 $R_{\odot}$	70 G
Radius of rotation	0.013 $R_{\odot}$
Maximum speed	30 km/s

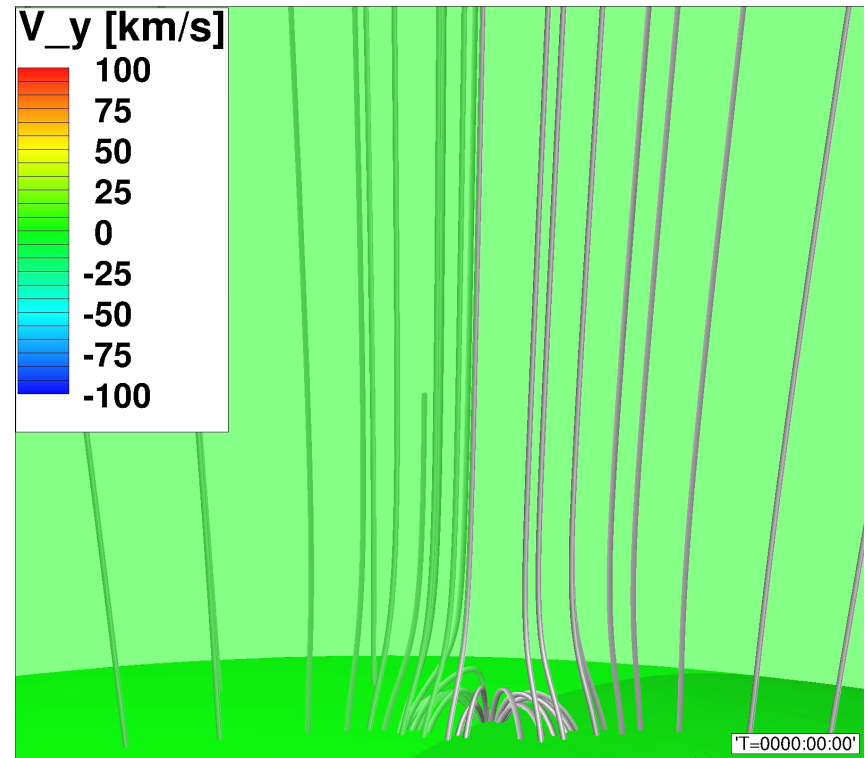


# Radial and out-of plane velocity profile across the center of the jet

Radial velocity with  
magnetic field

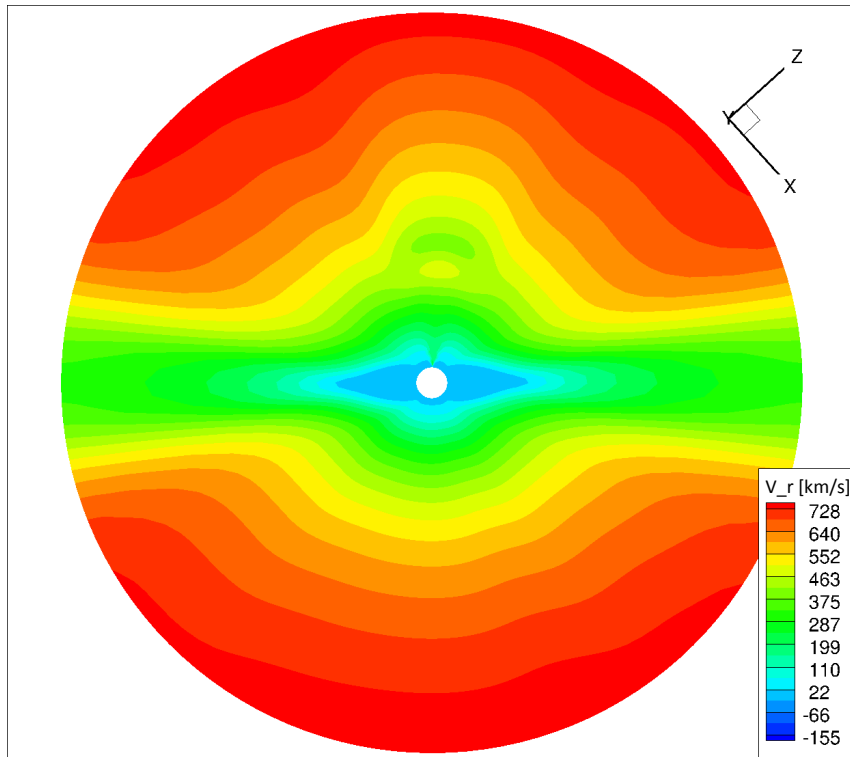


Out-of plane velocity with  
magnetic field

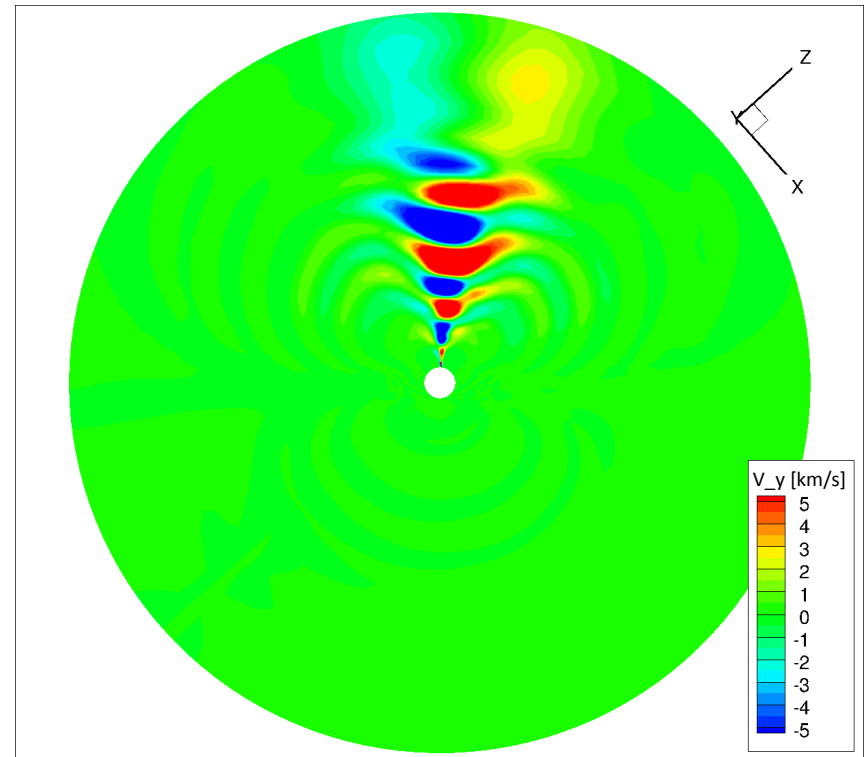


# Large scale velocity profiles ( $24 R_{\odot}$ ) ( $t = 11040$ s)

Radial velocity

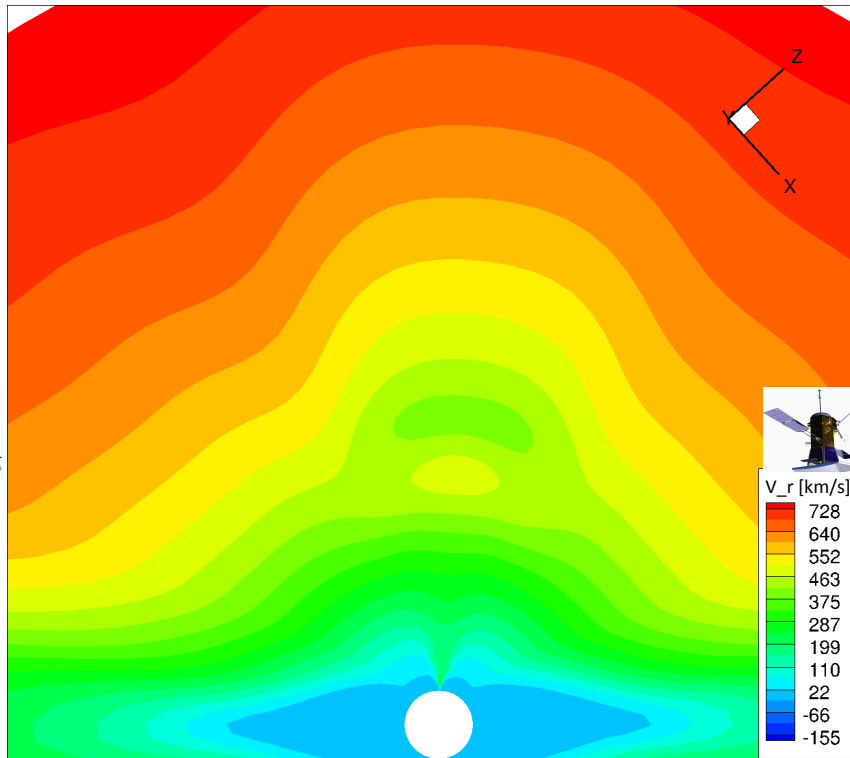


Out-of plane velocity

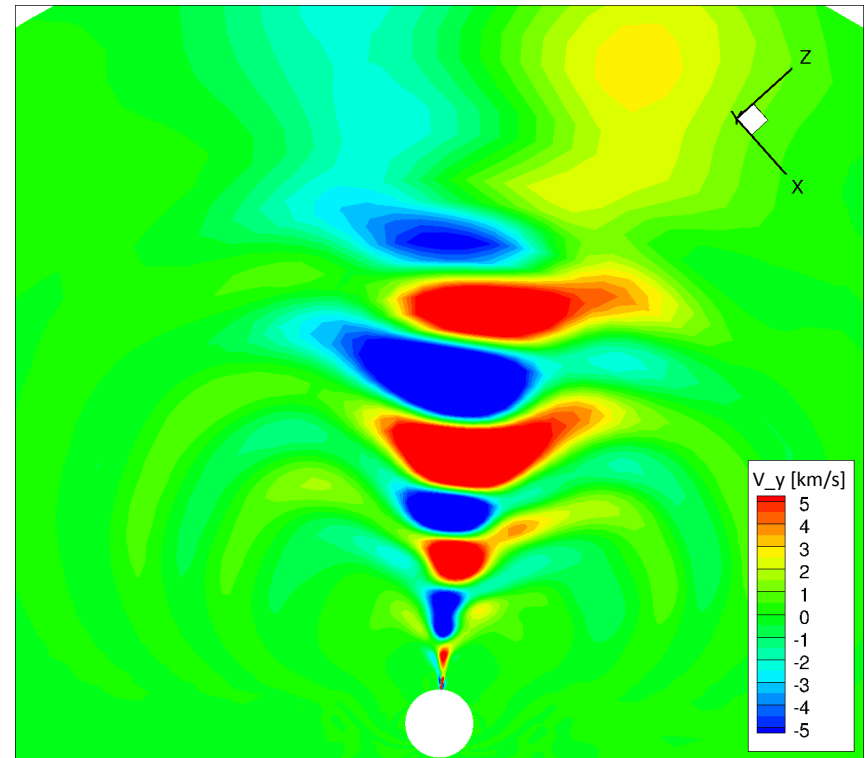


# Large scale velocity profiles ( $24 R_{\odot}$ ) ( $t = 11040$ s)

Radial velocity



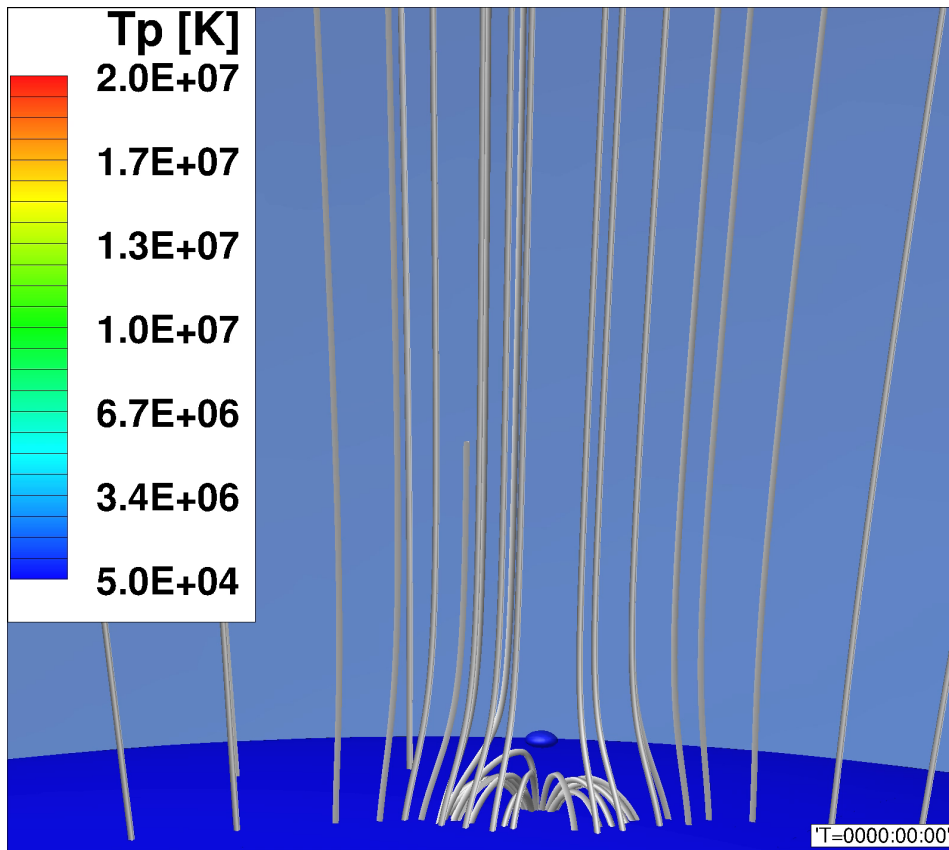
Out-of plane velocity



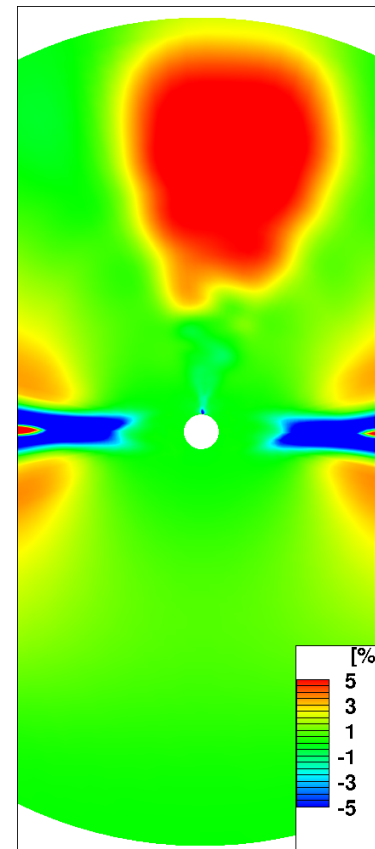
# Signatures in the magnetic field

## B=0 contour

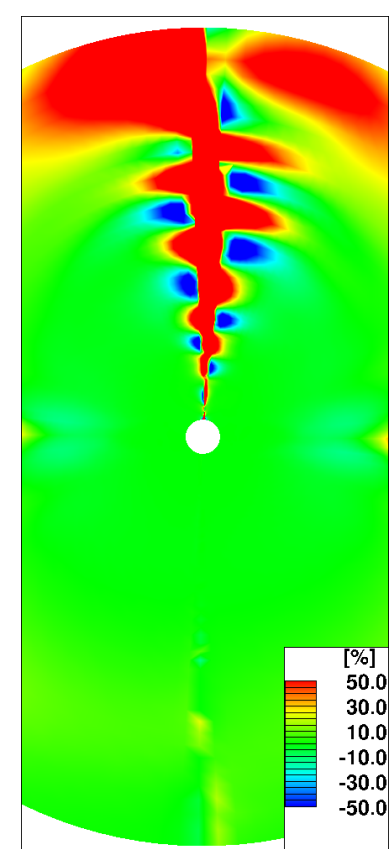
(color indicates proton temperature)



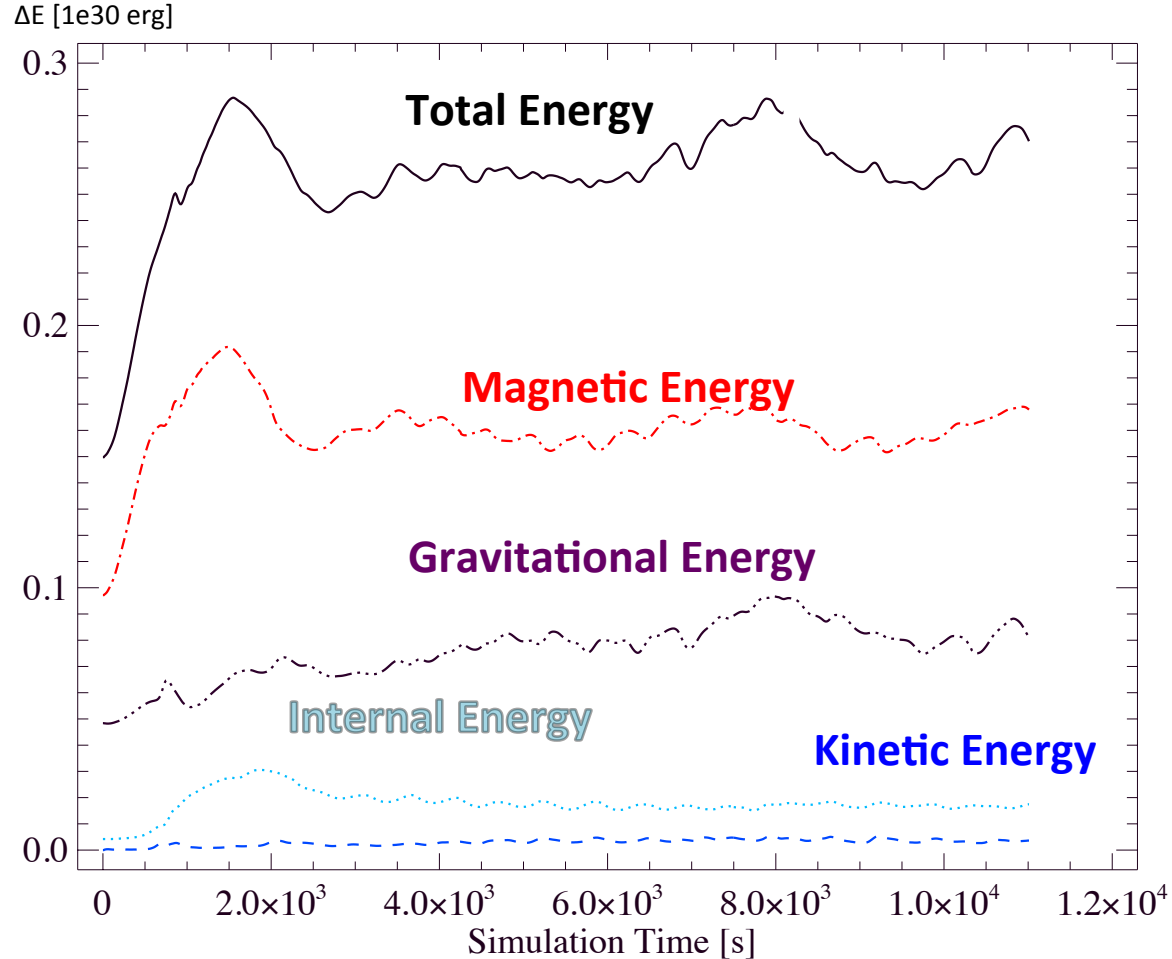
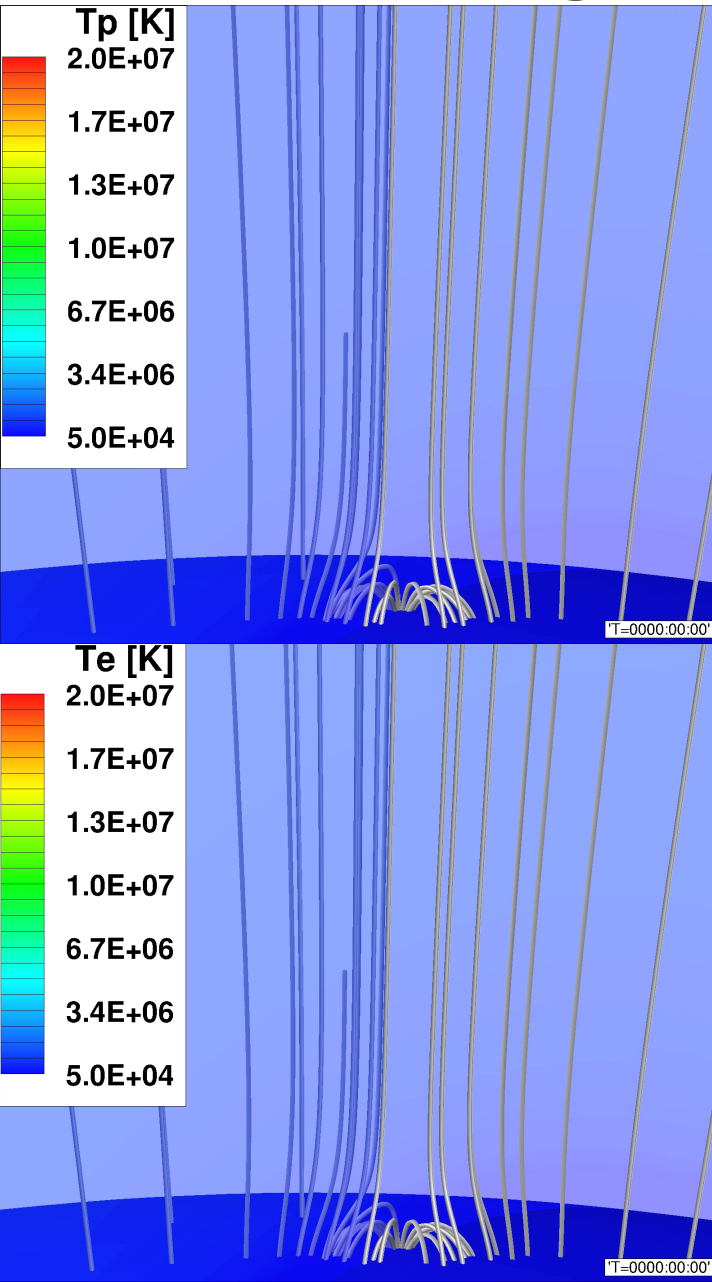
Relative  
change in  
magnetic  
energy



Relative  
change in  
azimuthal  
magnetic field

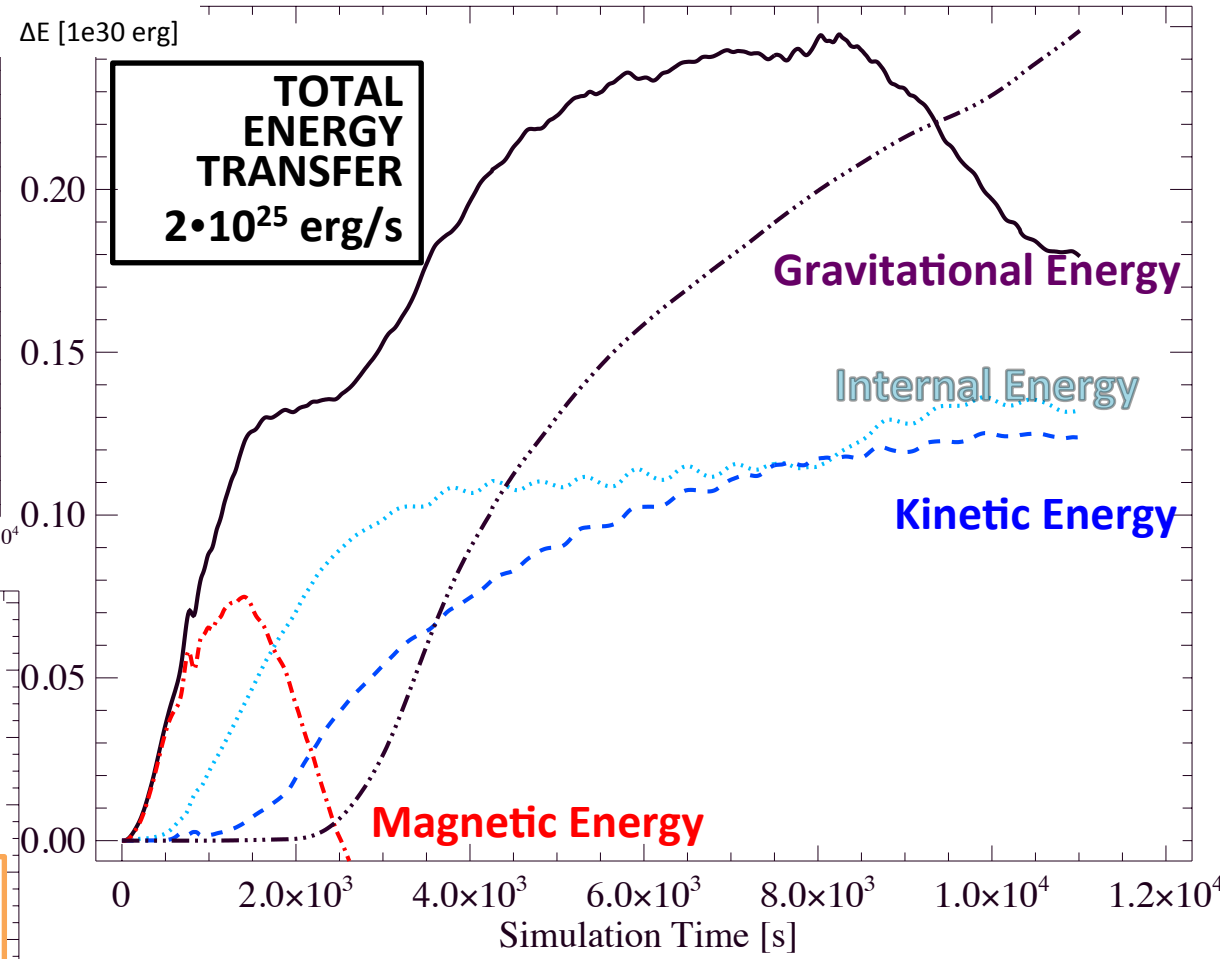
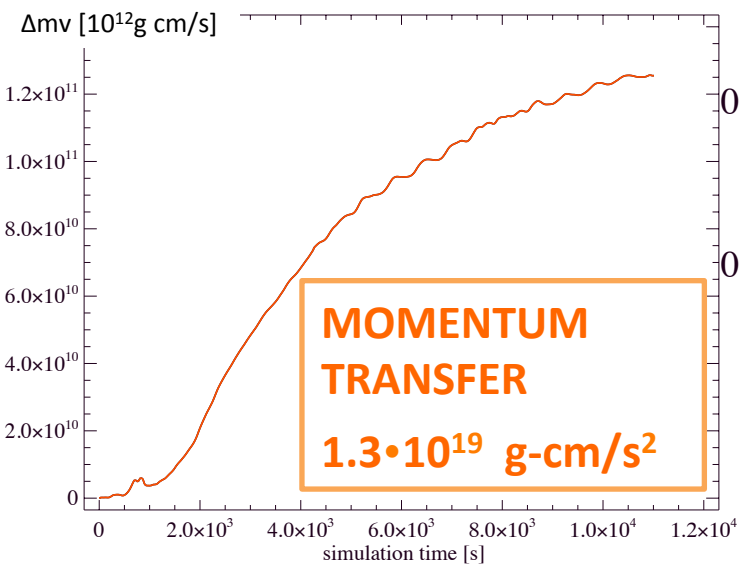
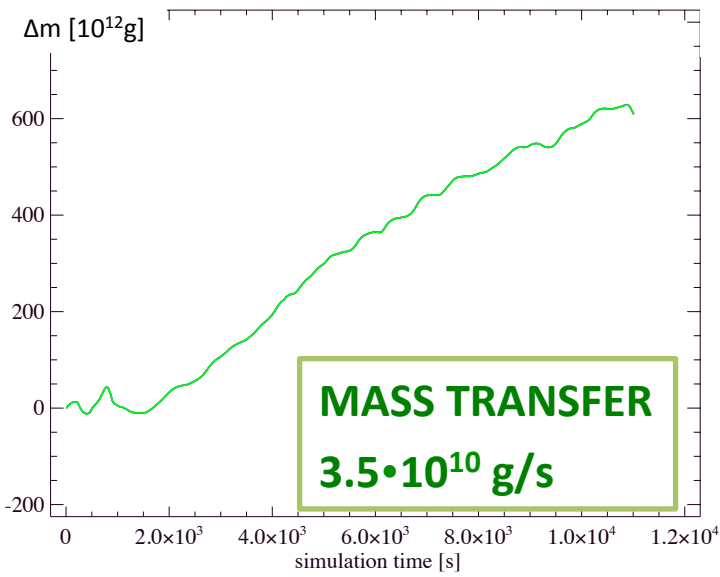


# Energetics in the jet's core



- the ion and electron temperatures show significant differences in the jet region
- magnetic energy is dominant in the core
- periodic changes

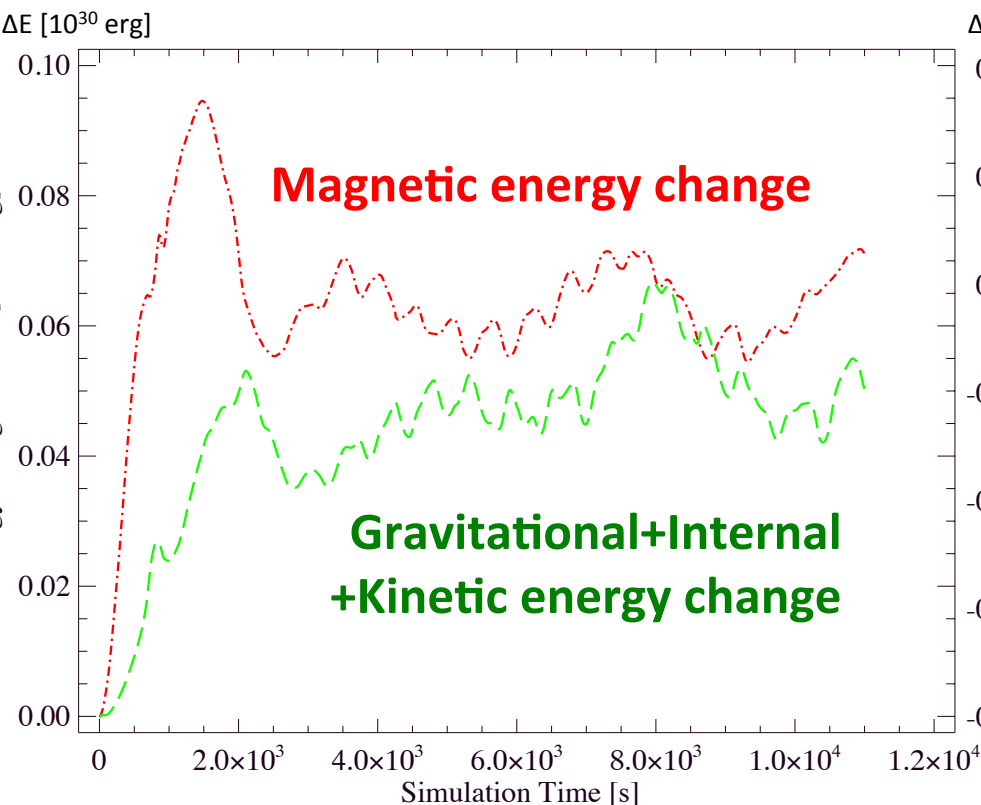
# Mass, momentum and energy transfer through the jet into the corona



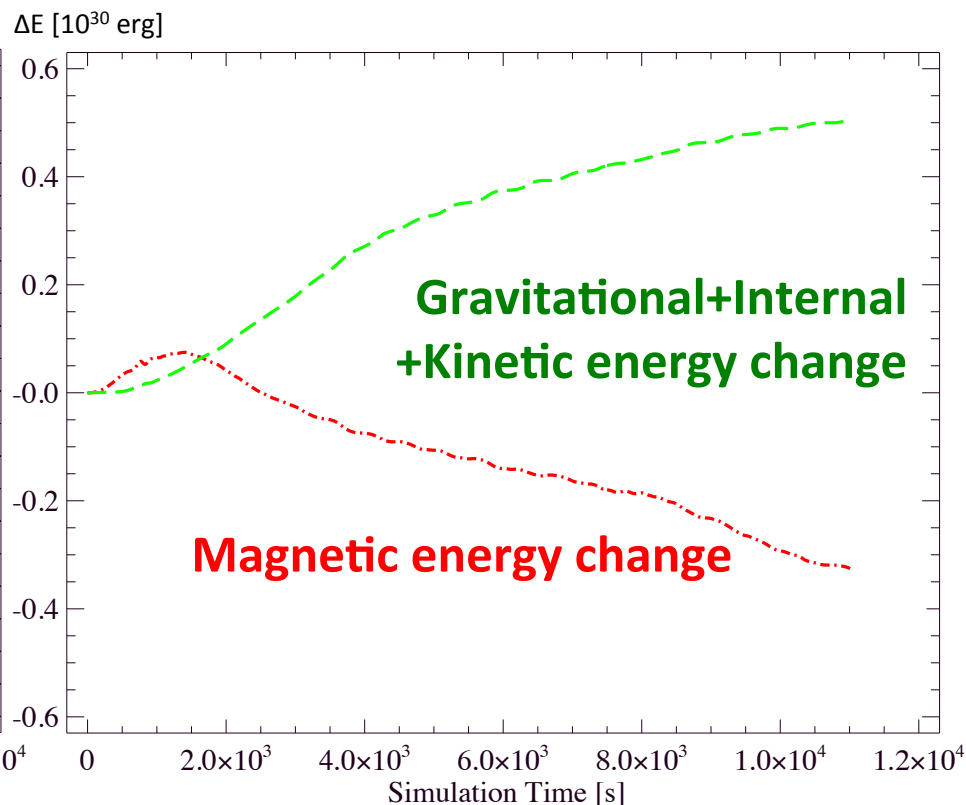
- magnetic energy is dominant only in the beginning
- gravitational energy change is dominant
- mass and momentum transport is significant compared to region with no jet

# Energetics of the jet

Local Scale:  
Jet core



Global Scale:  
Total Jet

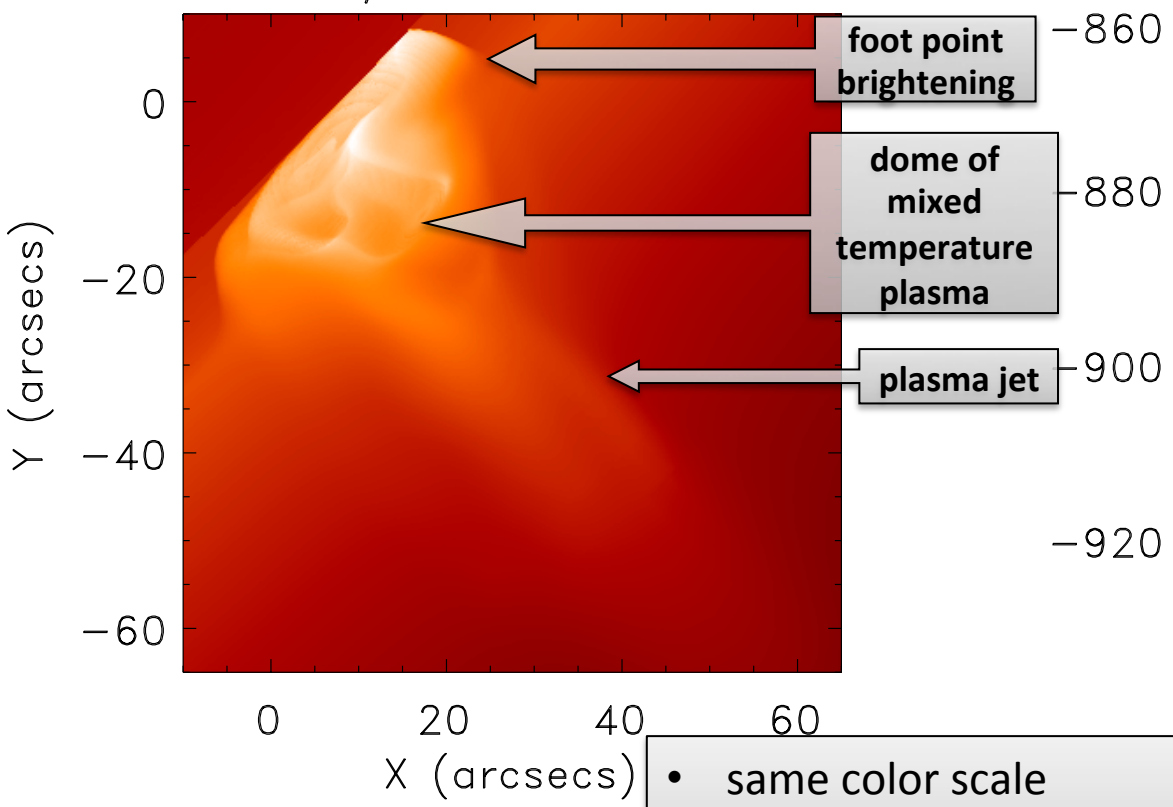




# Comparison of X-ray synthetic images to Hinode-XRT observations shows similarities

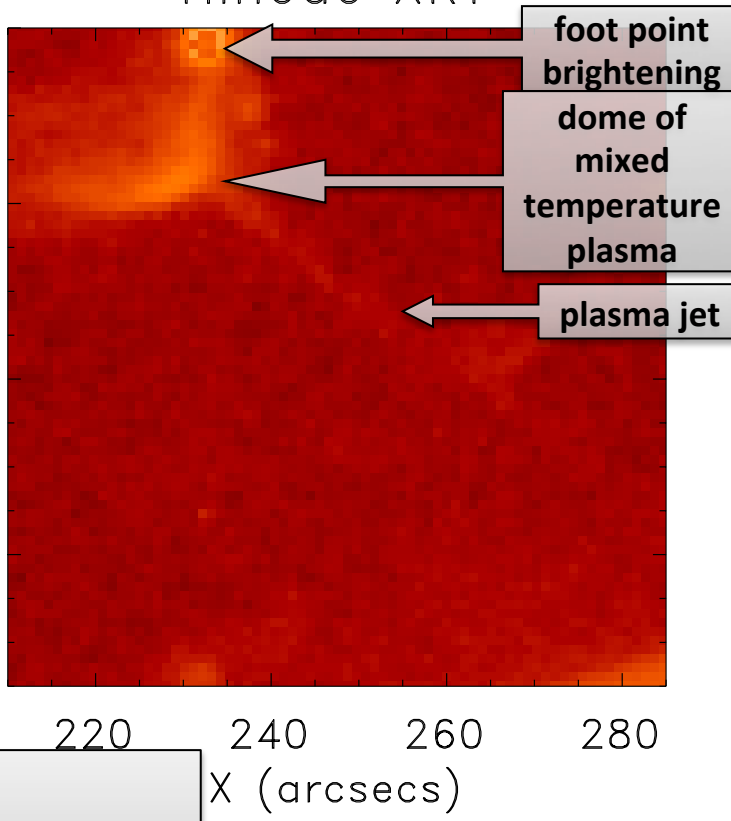
t = 01:03:30 simulation time

Synthetic XRT



2007-01-17T13:13:07

Hinode XRT



- same color scale
- same spatial scale
- also simulated AIA's EUV bands

# Conclusions

- The modeled jet compares well with observations.
- Inside perturbed corona up to  $24 R_{\odot}$  (relative to initial value)
  - **energy** transport:  $2 \cdot 10^{25}$  erg/s  
(12% increase in 3 hours)
  - **momentum** transport:  $1.3 \cdot 10^{19}$  g-cm/s<sup>2</sup>  
(102% increase in 3 hours)
  - **mass** transport:  $3.5 \cdot 10^{10}$  g/s  
(55% increase in 3 hours)
- Small local phenomena like a jet can produce global effects.
- We predict observable signatures for Solar Probe Plus.